Fermi/Large Area Telescope Discovery of Gamma-ray Emission from the Flat-spectrum Radio Quasar PKS 1454–354


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We report the discovery by the Large Area Telescope (LAT) onboard the Fermi Gamma-Ray Space Telescope of high-energy γ-ray (GeV) emission from the flat-spectrum radio quasar PKS 1454–354 (\( \gamma = 1.424 \)). On 2008 September 4, the source rose to a peak flux of \((3.5 \pm 0.7) \times 10^{-6} \) ph cm\(^{-2}\) s\(^{-1}\) (\( E > 100 \) MeV) on a timescale of hours and then slowly dropped over the following 2 days. No significant spectral changes occurred during the flare. Fermi/LAT observations also showed that PKS 1454–354 is the most probable counterpart of the unidentified EGRET source 3EG J1500–3509. Multiwavelength measurements performed during the following days (7 September with Swift; 6–7 September with the ground-based optical telescope Automated Telescope for Optical Monitoring; 13 September with the Australia Telescope Compact Array) resulted in radio, optical, UV, and X-ray fluxes greater than archival data, confirming the activity of PKS 1454–354.

**Key words:** galaxies: active – gamma rays: observations – quasars: individual (PKS 1454–354)

**Online-only material:** color figure

1. INTRODUCTION

PKS 1454–354 was discovered in the Molonglo Radio Telescope Survey at 408 MHz (Large et al. 1981) and later identified as a flat-spectrum radio quasar at \( z = 1.424 \) (Jackson et al. 2002; Hook et al. 2003). Very Long Baseline Array (VLBA) observations at 2.3 and 8.4 GHz showed a compact core with a weak one-sided jet (Petrov et al. 2005). Interest in this source was generated by the detection by EGRET onboard the Compton Gamma-Ray Observatory (CGRO) of the unidentified source 3EG J1500–3509 (Hartman et al. 1999). Mattox et al. (2001) tentatively indicated PKS 1454–354 as the only possible association with the EGRET unidentified source, but the association probability was very low (0.027). However, they noted that the EGRET probability contours were not closed on the eastern side and they suggested that this could be due to a contaminating source. The same possibility was proposed by Tornikoski et al. (2002) on the basis of millimeter observations, while Sowards-Emmert et al. (2004) suggested that PKS 1454–354 could be one of the two likely counterparts of the unidentified EGRET source (the other is PMJ J1505–3432).

The answer to this conundrum comes from the Fermi satellite and is presented here. The wide field of view (~2.4 sr) of the Large Area Telescope (LAT; Atwood et al. 2008) coupled with its high sensitivity (more than an order of magnitude better than EGRET) and spatial resolution (of the order of a few tens of arcminutes for \( E > 1 \) GeV) make it possible to monitor the whole sky every two orbits (~3 hr). During these scans, on 2008 September 4, a strong γ-ray flare was observed from the direction of PKS 1454–354 (Marelli 2008). Such alert triggered follow-up observations with Swift (Gehrels et al. 2004), performed on 7 September, with the Australia Telescope Compact Array (ATCA), done on 13 September and with the Automated Telescope for Optical Monitoring (ATOM), which started the observations on September 6.

Here, we report on the analysis of Fermi/LAT, Swift, ATCA, and ATOM data collected during the period of γ-ray activity and follow-up observations (Section 2) together with a comparison with archival data and a preliminary interpretation of the results (Section 3).

2. DATA ANALYSIS

2.1. Fermi/LAT

The LAT onboard the Fermi satellite detects photons with energy from 20 MeV to >300 GeV through the conversion to electron–positron pairs in a silicon tracker (Atwood et al. 2008).
The data have been analyzed using Science Tools version 9.7, a software package dedicated to the Fermi/LAT data analysis. Events that have the highest probability of being photons, labeled “diffuse” in the Science Tools, coming from zenith angles < 105° (to avoid Earth’s albedo) were selected. The diffuse emission from the Milky Way has been subtracted from the Third EGRET Catalog (Hartman et al. 1999), suggesting the “one month data”, with an upper limit of 2 × 10^{-8} ph cm^{-2} s^{-1} (90% confidence level, 0.2–300 GeV energy band) for PMN J1505–3432, the other candidate counterpart according to Sowards-Emmerd et al. (2004). PKS 1454–354 is detected in the one month data with a flux (4.5 ± 0.9) × 10^{-8} ph cm^{-2} s^{-1}, consistent with the flux of 3EG J1500–3509 reported in the Third EGRET Catalog (Hartman et al. 1999), suggesting that the FSRQ can be the likely counterpart and the distortion of the EGRET probability contours is due to some residual background or to some residual γ-ray emission from the source PMN J1505–3432 located at the boundaries of the EGRET source.

The LAT data of PKS 1454–354 were fitted with a power-law model for the spectrum (Figure 2). On 2008 September 4, the blazar was detected at a flux (1.29 ± 0.15) × 10^{-6} ph cm^{-2} s^{-1} in the 0.2–300 GeV energy band, with a peak of (1.7 ± 0.4) × 10^{-6} ph cm^{-2} s^{-1} between 06 and 12 UTC. Subsequently, the source was observed in a slower declining phase reaching a flux (0.56 ± 0.24) × 10^{-6} ph cm^{-2} s^{-1} after about 18–24 hr. In Figure 3 (upper panel), we present the corresponding γ-ray light curve integrated over 6 hr time bins between 3 September 00 UT and 2008 September 6 00 UT (MJD 54712–54715).

The photon index Γ, defined as \( F(E) \propto E^{-\Gamma} \), averaged over the day, is 2.0 ± 0.1 and it remains constant through the flare (see Figure 3, lower panel).

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61 http://fermi.gsfc.nasa.gov/ssc/data/analysis/software/
observation performed on 2008 January 1 at 00:16 UTC (ObsID 00036799001, exposure 9.0 ks) is available in the archive. The data from the hard X-ray detector (BAT; Barthelmy et al. 2005), the X-ray telescope (XRT; Burrows et al. 2005), and the optical/ultraviolet monitor (UVOT; Roming et al. 2005) were processed and analyzed with HEASOFT version 6.5 and the calibration database updated on 31 July 2008. The results are summarized in Tables 2 and 3.

The source is not detected by the BAT instrument, even after summing the two observations (exposure 11 ks). The upper limit (3σ) is 6 × 10^{-10} erg cm^{-2} s^{-1} in the 20–100 keV band. This is not surprising given the low exposure.

The XRT is operated in the photon-counting mode. Its data were processed with the xrtpipeline task using standard parameters and selecting single to quadruple pixel events (grades 0–12). The output spectra were rebinned to have at least 20 counts per energy bin. The X-ray spectrum measured on 7 September is best fitted (χ^2 = 1.3, degrees of freedom (dof) = 1) with a power-law model with a photon index Γ = 2.0 ± 0.5 with Galactic absorption (N_H = 6.46 × 10^{20} cm^{-2}; Kalberla et al. 2005). We note the very low statistics, which do not allow us to put tight constraints, particularly in the photon index. The flux in the 0.2–10 keV energy band was (2.2 ± 0.2) × 10^{-12} erg cm^{-2} s^{-1}, significantly higher than the value of (8.9 ± 0.9) × 10^{-13} erg cm^{-2} s^{-1} measured during the observation performed on 1 January, indicating that the blazar was still active 3 days after the γ-ray outburst detected by LAT. The photon indices of the two observations are, however, consistent within the measurement errors (Γ = 2.0 ± 0.3 on 1 January).

UVOT snapshots with all the six available filters, ν (5468 Å), b (4392 Å), u (3465 Å), uvw1 (2600 Å), uvw2 (2246 Å), uvw2 (1928 Å), taken during the obsID 00036799003 observation, were integrated with the uvotsource task and then analyzed by using the uvotimsum task, with a source region radius of 5′′ for the optical filters and 10′′ for the UV, while the background was extracted from an annular region centered on the blazar, with external radius of 25′′ and internal radius of 7′′ for the optical filters and 12′′ for the UV. It was not possible to select a larger external radius because of nearby sources. The observed magnitudes were m_v = 16.79 ± 0.09, m_b = 17.58 ± 0.06, m_u = 16.79 ± 0.05, m_uvw1 = 17.11 ± 0.08, m_uvw2 = 17.3 ± 0.2, m_uvw2 = 18.0 ± 0.1. These values were dereddened according to the extinction laws of Cardelli et al. (1989) with A_V = 0.34 and then converted into flux densities according to the standard formulae and zero points (Poole et al. 2008). During the observation of 1 January only one filter (uvw2) was available and the measured magnitude was 18.37 ± 0.05, about 1 mag greater than the value measured on 7 September.

2.3. Ground-Based Optical Observations

Optical observations in Johnson R and B filters were obtained on 6 September at 17:44 UTC and 7 September at 18:30 UTC with the 0.8 m optical telescope ATOM in Namibia. ATOM is operated robotically and obtains automatic observations of confirmed or potential gamma-bright blazars. Data analysis (debiasing, flat fielding, photometry using SExtractor; Bertin & Arnouts, 1996) is conducted automatically. PKS 1454–354 does not appear to be spatially extended beyond the point-spread function and no corrections for potential contributions from the host galaxy were performed. Absolute photometry was obtained using nonvariable field stars in photometric nights using several calibration stars. The observed magnitudes are

Fitting the data with a broken power-law model does not improve the fit quality implying that no break is detected in the energy range 0.2–5 GeV (see Figure 2). The results of the LAT analysis are summarized in Table 1. The quoted errors are just statistical ones. The systematic errors should be added. According to the studies on the Vela Pulsar (Abdo et al. 2008), our current conservative estimates of systematic errors are ≲30% for flux measurements and 0.1 for the photon index. Significant reduction of such systematic uncertainties is expected once the calibration of the LAT instrument will be completed.

The results of the LAT spectral analysis were confirmed with two additional independent methods: the first one uses a spectrum unfolding based on the Bayes’ theorem (D’Agostini 1995), while the second one uses the nested HEALpix (Górski et al. 2005).

2.2. Swift

PKS 1454–354 was observed by Swift on 2008 September 7 at 11:50 UTC (ObsID 00036799003, exposure 1.5 ks). Another
and north–south spur, providing instantaneous two-dimensional 
inner five antennas were located on the main east–west track 
of 90 m and the sixth antenna located 4 □. 
with a maximum separation between the inner five antennas 
with the ATCA at 2 □.

 Shortly after detecting the high γ-ray flux reported in this 
paper, on 2008 September 13, PKS 1454 □ was observed 
with the ATCA at 2.368, 4.800, 8.640, 18.496, and 19.520 
GHz. At this time the array was in its H75 configuration, 
with a maximum separation between the inner five antennas 
of 90 m and the sixth antenna located 4.3 km to the west. 
The inner five antennas were located on the main east–west track 
and north–south spur, providing instantaneous two-dimensional 
(u, v) coverage.

### 2.4. Radio Observations

The ATCA primary calibrator, PKS 1934-638, was observed 
first and used to calibrate the array at all frequencies over the 128 
MHz observing bandwidth. The data were reduced following the 
standard procedures in MIRIAD, and the task uvfilux used 
to fit a point source model to the data and determine the flux 
density and rms scatter about this point. Inspection of the data 
confirmed that the point source approximation was valid at all 
frequencies. Phase stability was poor on the long baselines to 
so this antenna was omitted from the analysis. The fitted flux 
used standard formulae and zero points. Comparison 
with Swift/UVOT data indicates some intraday variability.

### Table 2

Summary of the Swift Data

<table>
<thead>
<tr>
<th>ObsID</th>
<th>Exposure (ks)</th>
<th>$N_H$ $(10^{20}$ cm$^{-2}$)</th>
<th>$\Gamma$</th>
<th>Flux $0.2$–$10$ keV $(10^{-12}$ erg cm$^{-2}$ s$^{-1}$)</th>
<th>$\chi^2$/dof</th>
</tr>
</thead>
<tbody>
<tr>
<td>BAT</td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>00036799001/3 (sum)</td>
<td>11</td>
<td>6.46</td>
<td>2.0 ± 0.3</td>
<td>0.89 ± 0.09</td>
<td>8.0/6</td>
</tr>
<tr>
<td>XRT</td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>00036799001</td>
<td>9.0</td>
<td>6.46</td>
<td>2.0 ± 0.3</td>
<td>0.89 ± 0.09</td>
<td>8.0/6</td>
</tr>
<tr>
<td>00036799003</td>
<td>1.5</td>
<td>6.46</td>
<td>2.0 ± 0.5</td>
<td>2.2 ± 0.2</td>
<td>1.3/1</td>
</tr>
</tbody>
</table>

**Note.** See the text for details.

### Table 3

Summary of the Swift Data (UVOT)

<table>
<thead>
<tr>
<th>ObsID</th>
<th>$\nu$ (5468 Å)</th>
<th>$\nu$ (4392 Å)</th>
<th>$\nu$ (3465 Å)</th>
<th>$\nu$ (2600 Å)</th>
<th>$\nu$ (2246 Å)</th>
<th>$\nu$ (1928 Å)</th>
</tr>
</thead>
<tbody>
<tr>
<td>00036799001</td>
<td>18.37 ± 0.05</td>
<td>17.3 ± 0.2</td>
<td>18.0 ± 0.1</td>
<td>17.3 ± 0.2</td>
<td>17.3 ± 0.2</td>
<td>18.0 ± 0.1</td>
</tr>
<tr>
<td>000367999001</td>
<td>16.79 ± 0.09</td>
<td>16.79 ± 0.05</td>
<td>17.11 ± 0.08</td>
<td>17.1 ± 0.08</td>
<td>17.1 ± 0.08</td>
<td>17.1 ± 0.08</td>
</tr>
</tbody>
</table>

**Note.** See the text for details.

$B = 17.01 ± 0.02$ and $R = 16.46 ± 0.02$ on 6 September, 
and $B = 16.83 ± 0.02$ and $R = 16.53 ± 0.02$ on 7 September. 
The magnitudes were then corrected for galactic extinction 
using the same methods outlined above for Swift/UVOT and converted 
in fluxes using standard formulae and zero points. Comparison 
with Swift/UVOT data indicates some intraday variability.
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Figure 4. VLBI images of PKS 1454–354 at 8.4 GHz. The left panel shows the naturally weighted image and the right panel the tapered image obtained from imaging the final calibrated data with down-weighted visibilities of long baselines (10% at 100 million wavelengths) to increase the signal-to-noise ratio of the low-surface brightness jet. The jet extends to ~50 mas to the west–northwest. Both images have an rms noise of ~0.08 mJy beam$^{-1}$ and the lowest contour is at 3.7 and 3.8 times the rms noise, respectively. The restoring beam is shown on bottom left of each image (2.87 × 0.75 mas at 4; 4.9 × 3.7 mas at ~60°). These data were obtained on 2007 November 10, as part of the TANAMI program (tracking AGN with Austral Milliarcsecond Interferometry; Ojha et al. 2008) with six elements of the Australian South-African LBA (Long Baseline Array). See the text for details.

Figure 5. SED of PKS 1454–354. ATCA (13 September), ATOM (6 and 7 September), Swift (XRT and UVOT; 7 September), and Fermi (LAT; 4 September) data are indicated with filled circles (red) and connected by a dashed line for a better visualization. Fermi/LAT data of August and Swift (XRT and UVOT) observation of 2008 January 1, to which we refer as "low state," are indicated with squares (green). Archival data are marked with filled triangles (black). Radio data: Molonglo, 408 MHz (Large et al. 1981); NVSS, 1.4 GHz (Condon et al. 1998); Parkes, 4.85 GHz (Wright et al. 1996); CRATES, 8.4 GHz (Healey et al. 2007); ATCA, 20 GHz (Massardi et al. 2008); SUMSS, 843 MHz (Mauch et al. 2003); VCS3, 2.3 and 8.6 GHz (Petrov et al. 2005). Optical: USNO B1, B, R, I filters (Monet et al. 2003); Guide Star Catalog II, B and R filters (Lasker et al. 2008). X-rays: ROSAT Faint Source Catalog, 1 keV (Voges et al. 2000). γ-Rays: EGRET, 400 MeV (Hartman et al. 1999).

(A color version of this figure is available in the online journal.)

at 19.904 GHz of 0.90±0.05 Jy determined from an observation in May 2007, indicating the source had brightened by 40% by September 2008. The PMN survey (Wright et al. 1996) recorded a flux density of 0.566 ± 0.041 Jy at 4.85 GHz in November 1990. It can be concluded that, at the time of the γ-ray outburst, PKS 1454–354 was in a comparatively bright radio state showing an inverted spectrum, which can be interpreted as the indication of an enhanced ongoing radio activity. Single-dish radio flux-density monitoring observations are ongoing at the Ceduna 30 m telescope and will be reported elsewhere.

Figure 4 shows very long baseline interferometry (VLBI) images of PKS 1454–354 at a frequency of 8.4 GHz. These data were obtained on 2007 November 10, as part of the TANAMI program (tracking active galactic nucleus (AGN) with Austral Milliarcsecond Interferometry; Ojha et al. 2008) with the Australian Long Baseline Array consisting of telescopes at Parkes, Narrabri, Mopra, Hobart, Ceduna (all in Australia), and Hartbeesthoek (South Africa). The data were correlated using the DiFX software correlator (Deller et al. 2007) at the Swinburne University of Technology. Details of these telescopes as well as the data calibration path are described in Ojha et al. (2005). The approximate position angle of the diffuse jet is ~300°. The correlated flux density of the milliarcsecond-scale core–jet structure is about a factor of 2 below the (nonsimultaneous) total flux density at 8.64 GHz measured by ATCA in September 2008: ~415 mJy, of which more than 95% is contained within the compact core, i.e., on scales smaller than 0.75 mas (the full width at half-maximum of the minor axis of the beam). Model fitting the core with an elliptical Gaussian function results in a model component with 400 mJy and a size of 0.26 × 0.17 mas, corresponding to a brightness temperature limit of $T_B = 1.6 	imes 10^{11}$ K. These are the first published VLBI images of PKS 1454–354 at this angular resolution and dynamic range. The only other radio structural information on these scales comes from the VCS3 program (Petrov et al. 2005) carried out in May 2004 and showing no structure at 8.4 GHz and a jet to the west–northwest at 2.3 GHz consistent with the morphology that we find at 8.4 GHz. Petrov et al. (2005) report 0.51 Jy at both 2.3 and 8.4 GHz, and unresolved components of 0.38 and 0.21 Jy, respectively.

3. DISCUSSION OF THE SED

In Figure 5, we report the spectral energy distribution (SED) built with archival data together with Fermi/LAT, Swift, ATCA, and ATOM data analyzed in the present work. Two different states are shown for γ-rays (LAT): the quiescence value, measured from the one month data set, and the active state, obtained from the data collected on 4 September, the day of the discovery. A change in flux during the flare of more than 1 order of magnitude is evident.

Fluxes at X-rays, ultraviolet, optical, and radio frequencies are higher than those of the archival data, although the Swift...
observation was performed on 7 September and ATCA measurements were done on 13 September. This confirms that PKS 1454–354 was still active at all the wavelengths several days after the γ-ray flare observed by LAT. It is worth noting the absence of significant spectral changes in the X- and γ-ray bands, but given the low statistics we cannot exclude the presence of spectral variations on short timescales. However, at radio frequencies, there was an inversion of the spectral index $\alpha_r$, defined as $S_r \propto \nu^{-\alpha_r}$. The value calculated from archival data is positive ($\alpha_r^{6.4 \, \text{GHz}} = 0.040$, Sowards-Emmerd et al. 2004; $\alpha_r^{8.6 \, \text{GHz}} = 0.054$, Healey et al. 2007; with an error of about 7% in both cases) and becomes negative during the outburst ($\alpha_r^{2.3 \, \text{GHz}} = -0.10 \pm 0.01$). This can be interpreted as the signature of the emission of a relativistic plasma blob.

In order to investigate the SED characteristics and compare with larger samples of similar objects, two-point spectral indices have been calculated according to the formula of Ledden & O’Dell (1985). We used radio data at 4.8 GHz, optical data at $\lambda = 3465$ Å (Swift/UVOT $\nu$ filter, to minimize the contribution of the host galaxy), and X-rays at 1 keV. From ATCA radio measurements of 13 September, optical/UV observations from Swift/UVOT and X-rays from Swift/XRT collected on 7 September, we calculated $\alpha_{ox} = -0.1$, $\alpha_{ro} = 1.85$, and $\alpha_{rx} = 1.0$, respectively. Then, the flux densities have been $K$-corrected by the multiplicative factor $(1 + z)^{1-\alpha_r}$. Finally, we obtain the following values: $\alpha_{o\delta} = 1.11$, $\alpha_{o\nu} = 0.49$, $\alpha_{r\delta} = 0.69$.

These values, particularly $\alpha_{r\delta} < 0.5$, seem to suggest that this FSRQ belongs to the class of high-frequency peaked FSRQ (HFSRQ) conjectured by Padovani et al. (2003), which have the synchrotron emission peaked at UV/soft X-rays. However, the shape of the SED, specifically the strong change in spectral index between optical and X-ray frequencies ($\alpha_{ro} = 1.85$ and $\alpha_{rx} = 1.0$), clearly rules out this possibility.

Instead, our two-point spectral indices are very similar to those of X-ray-selected FSRQ analyzed in Maraschi et al. (2008), where, in addition to the usual nonthermal model of blazars (synchrotron self-Compton, SSC, plus external Compton, EC), there is a component due to the X-ray corona above the accretion disk (for the disk–jet connection in FSRQ see also Grandi & Palumbo 2004; Sambruna et al. 2007). The moderately high value of $\alpha_{o\delta}$, together with a soft X-ray spectrum ($\alpha_{rx} = 1$), suggests the presence of a corona, which is best observed when the source is in quiescence and the jet contribution is low. The contrary occurs during γ-ray flares, due to the dominance of the jet emission over the corona. In this case, the X-ray spectrum should become harder ($\alpha_{r} < 1$), but the measured X-ray photon index has a large error ($\Gamma = 2.0 \pm 0.5$) and, although consistent with this hypothesis, it is not conclusive.

To support this hypothesis, we calculate some order-of-magnitude values for the physical parameters and we refer to Maraschi et al. (2008) for more details. These authors found that $\alpha_{o\delta}$ is correlated with the viewing angle and, therefore, is anticorrelated with the Doppler factor ($\delta$). Taking advantage of such correlation, we can estimate a value of $\delta = 15–20$, which in turn allows us to calculate the size of the emitting region in the source frame $R'$. According to the well-known formula $R' \sim c \delta \Gamma / (1 + \delta z)$, the emitting region must be smaller than $R' \lesssim 10^{16}$ cm during the rise of the γ-ray flare, whose timescale is estimated to be $\approx 12$ hr (see Figure 3). The longer timescale during the declining part of the flare suggests a larger emitting region, meaning that the electrons escaped from the processing region, resulting in a decrease of the jet power. A similar effect can be obtained decreasing the density of seed photons, but in this case we expected a change in the optical emission too, disentangled from the synchrotron emission. However, the coordinated increase of the whole synchrotron hump, from radio to optical frequencies, makes the change in the injected power the most likely explanation for the observed variability.

The jet-frame radiation energy density can be estimated by the energy density of the broad-line region (BLR) according to the formula $U_{\text{BLR}} = U_{\text{BLR}}(85^2$ (given that the Doppler factor is almost equal to the bulk Lorentz factor). Since the BLR luminosity is assumed to be 10% of $L_{\text{disk}}$ ($\sim 4 \times 10^{46}$ erg s$^{-1}$ in our case), the size of BLR can be calculated as $R_{\text{BLR}} \sim 10^{17} \sqrt{L_{\text{disk}} / 4 \pi c K_{\text{BLR}}} = 6 \times 10^{17}$ cm, where $L_{\text{disk}} = 4 \pi c K_{\text{BLR}}$ is the disk luminosity in units of $10^{45}$ erg s$^{-1}$. Then, the BLR energy density is $U_{\text{BLR}} = L_{\text{BLR}} / (4 \pi c K_{\text{BLR}}) = 0.029$ erg cm$^{-3}$. The corresponding total radiation energy density in the jet frame is therefore $U_{\text{rad}} \sim 6–12$ erg cm$^{-3}$, depending on the value of $\delta$, which, in turn, results in a peak electron Lorentz factor of the order of $\gamma_{\text{peak}} \sim 100–1000$. These values are close to the parameters obtained with a more detailed modeling of similar FSRQ reported in Maraschi et al. (2008).

The large errors, particularly in the X-ray photon index, leave open different possibilities, involving one or two SSC components (e.g., Ballo et al. 2002; Giommi et al. 2008). In addition, flat X-ray spectra from blazar jets may possibly be a signature of pair cascades (e.g., Svensson 1987; Mücke et al. 2003; Böttcher et al. 2008).

4. FINAL REMARKS

We have reported the discovery by Fermi/LAT of γ-ray emission from the FSRQ PKS 1454–354. LAT data showed flux variability on 6–12 hr timescale, but with negligible spectral changes. The blazar is positionally consistent with the unidentified source 3EG J1500–3509 and, therefore, it is likely to be the EGRET source counterpart. The increased activity measured at radio, optical, ultraviolet, and X-ray frequencies in the days following the γ-ray outburst strengthens the identification of PKS 1454–354 as a high-energy γ-ray source.

The analysis of the SED and some order-of-magnitude calculations suggest that PKS 1454–354 is a typical FSRQ, although the large measurement errors hampered our capability to put tight constraints on the origin of the X-ray emission: the soft X-ray spectrum can be explained either in terms of SSC or of a corona above the accretion disk, while the high-energy peak is likely to be due to EC processes. However, other theories (e.g., hadronic models) can also be considered to explain the observed data.

Further, more precise multiwavelength observations of PKS 1454–354 and, specifically, a spectral coverage at submillimeter, IR, and optical frequencies, where the synchrotron peak is expected, will shed further light on the modeling of this source, allowing a more precise correlation between the synchrotron and γ-ray inverse-Compton peak in quiet and flaring states, clarifying the astrophysical mechanisms at work in this source.

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